

Alexandre Toubiana



# Introduction to model selection and hierarchical inference

PhD course, Milano-Bicocca, 15/05/2026

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# Plan of the lecture

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- ❖ Bayesian model selection
- ❖ Hierarchical inference
- ❖ Examples of astrophysical inference
- ❖ Posterior predictive checking

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# Bayesian hypothesis testing

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- ❖ The denominator in Bayes' Theorem

$$p(\theta|x, M) = \frac{p(x|\theta, M)p(\theta|M)}{p(x|M)}$$

is the **Bayesian evidence**

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Can be used for model selection  
**Posterior odds ratio:**

$$\mathcal{O}_{12} = \frac{p(x|M_1)p(M_1)}{p(x|M_2)p(M_2)}$$

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**Bayes Factor**

**Prior odds ratio.**

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- ❖ Gaussian likelihood:

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"Occam's razor"

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- ❖ After  $N_{\text{obs}}$  observations:

$$B_{12}^{(N_{\text{obs}})} \approx \frac{2\Delta\sqrt{N_{\text{obs}}}}{\sqrt{2\pi}\sigma} \exp \left( -\frac{N_{\text{obs}}\bar{x}^2}{2\sigma^2} \right).$$

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# Interpreting Bayes' factors

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- ❖ Interpretation is somewhat arbitrary
- ❖ Kass and Raftery (1995) suggested the following scale:

Bayes Factor	Interpretation
$< 3$	No evidence of $M_1$ over $M_2$
$> 3$	Positive evidence for $M_1$
$> 20$	Strong evidence for $M_1$
$> 150$	Very strong evidence for $M_1$

Correspond to “p-values” of 0.25, 0.05, 0.007

- ❖ In practice, posterior odds ratios can also be used as a test statistic, with significance and power computed via simulation in the usual (frequentist) way.

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# Computing Bayes' factors

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❖ **Direct integration:**

$$p(x|M) = \int p(x|\theta, M)p(\theta|M)d\theta \simeq \frac{1}{N} \sum_{\theta_i \sim p(\theta|M)} p(x|\theta_i, M)$$

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- ❖ **Other techniques:**
  - thermodynamic integration
  - nested sampling
  - reversible-jump MCMC

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# Hierarchical Models

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**based on Mandel, Farr, Gair MNRAS 2018,  
Taylor, Gerosa, PRD 2018**

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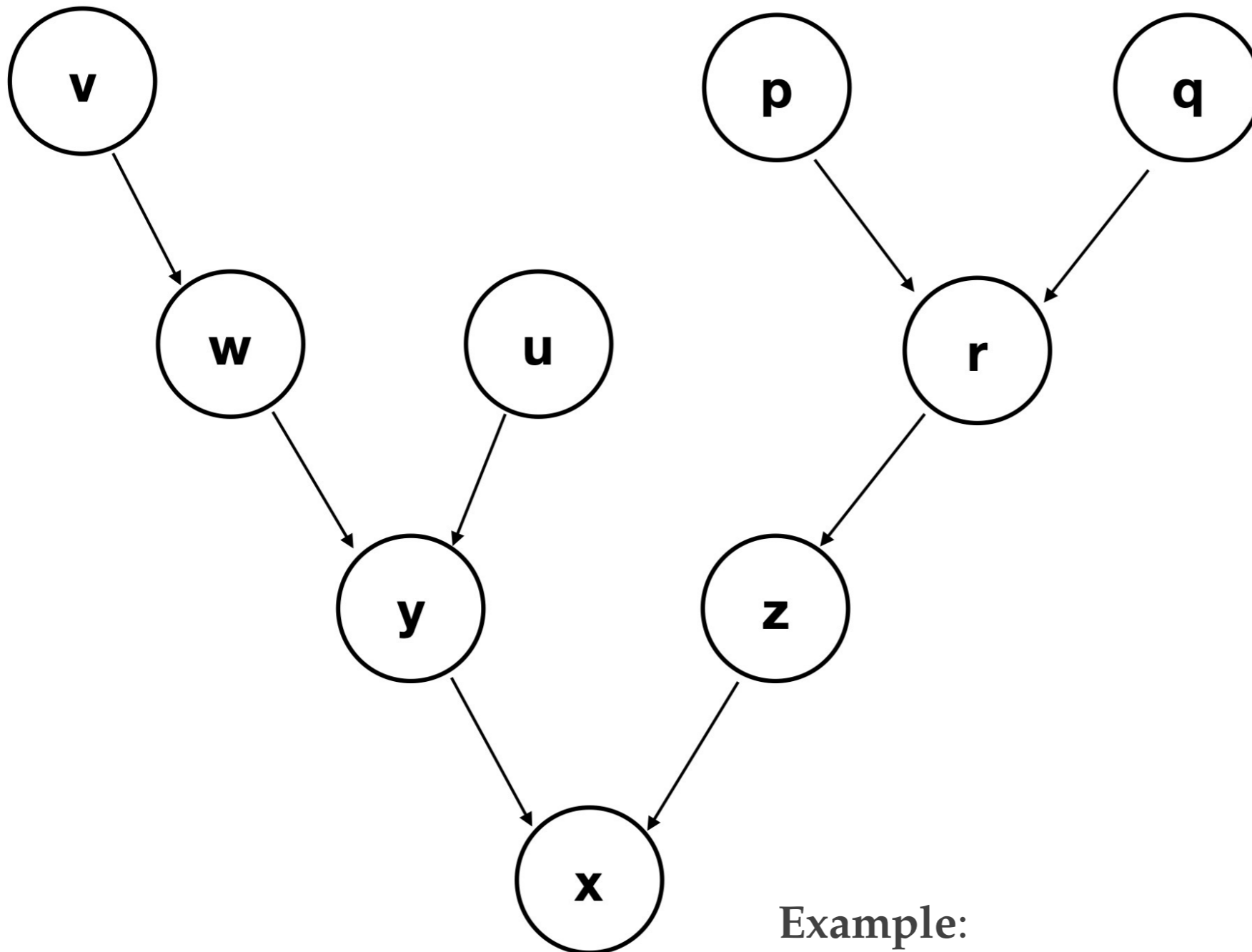
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**“Continuous model selection”**

# Graphical Model



Example:

$$p(p, q, r, s, t, u, v, w, x, y, z) = p(x | y, z) p(y | u, w) p(z | r) p(w | v) p(r | p, q) p(v) p(u) p(p) p(q)$$

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# Gaussian example

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- ❖ Population prior:

$$p(\theta|\Lambda) = \frac{1}{\sqrt{2\pi\sigma_{\text{pop}}^2}} \exp\left[-\frac{(\theta - \mu)^2}{2\sigma_{\text{pop}}^2}\right] \quad \Lambda = (\mu, \sigma_{\text{pop}}^2)$$

- ❖ Single-event likelihood:

$$p(x|\theta) = \frac{1}{\sqrt{2\pi\sigma_{\text{PE}}^2}} \exp\left[-\frac{(x - \theta)^2}{2\sigma_{\text{PE}}^2}\right]$$

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- ❖ Population likelihood:

$$p(x|\Lambda) = \frac{1}{\sqrt{2\pi(\sigma_{\text{PE}}^2 + \sigma_{\text{pop}}^2)}} \exp\left[-\frac{(x - \mu)^2}{2(\sigma_{\text{PE}}^2 + \sigma_{\text{pop}}^2)}\right]$$

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❖ Combining events:

$$p(\{x\}|\Lambda) = \frac{1}{(2\pi(\sigma_{\text{PE}}^2 + \sigma_{\text{pop}}^2))^{N_{\text{obs}}/2}} \exp\left[-\frac{N_{\text{obs}}(\bar{x} - \mu)^2}{2(\sigma_{\text{PE}}^2 + \sigma_{\text{pop}}^2)}\right] \exp\left[-\frac{N_{\text{obs}}\bar{s}^2}{2(\sigma_{\text{PE}}^2 + \sigma_{\text{pop}}^2)}\right]$$

$$\bar{x} = \frac{1}{N_{\text{obs}}} \sum_{i=1}^{N_{\text{obs}}} x_i \quad \bar{s}^2 = \frac{1}{N_{\text{obs}}} \sum_{i=1}^{N_{\text{obs}}} (x_i - \bar{x})^2$$

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- ❖ Maximum likelihood:  $\mu = \bar{x} \quad \sigma_{\text{pop}}^2 = \bar{s}^2 - \sigma_{\text{PE}}^2$
- ❖ Expanding around maximum in the large  $N_{\text{obs}}$  limit (“Fisher approximation”)  
Assuming broad priors:

$$p(\mu, \sigma_{\text{pop}}^2|\{x\}) = \frac{N_{\text{obs}}}{2\sqrt{2\pi}\bar{s}^3} \exp\left[-\frac{N_{\text{obs}}}{2} \left( \frac{(\mu - \bar{x})^2}{\bar{s}^2} + \frac{(\sigma_{\text{pop}}^2 - (\bar{s}^2 - \sigma_{\text{PE}}^2))^2}{2\bar{s}^4} \right)\right]$$

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$$\xi(\Lambda) = \int p(\text{det}|\theta) p(\theta|\Lambda) d\theta \quad p(\text{det}|\theta) = \int_{F(x) > \text{threshold}} p(x|\theta) dx.$$

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$$p(\{x\}|\Lambda) = \frac{N_{\text{det}}(\Lambda)^{N_{\text{obs}}}}{\xi(\Lambda)^{N_{\text{obs}}}} e^{-N_{\text{det}}(\Lambda)} \prod_i \int p(x_i|\theta) p(\theta|\Lambda) d\theta$$
$$= N(\Lambda)^{N_{\text{obs}}} e^{-N_{\text{det}}(\Lambda)} \prod_i \int p(x_i|\theta) p(\theta|\Lambda) d\theta$$

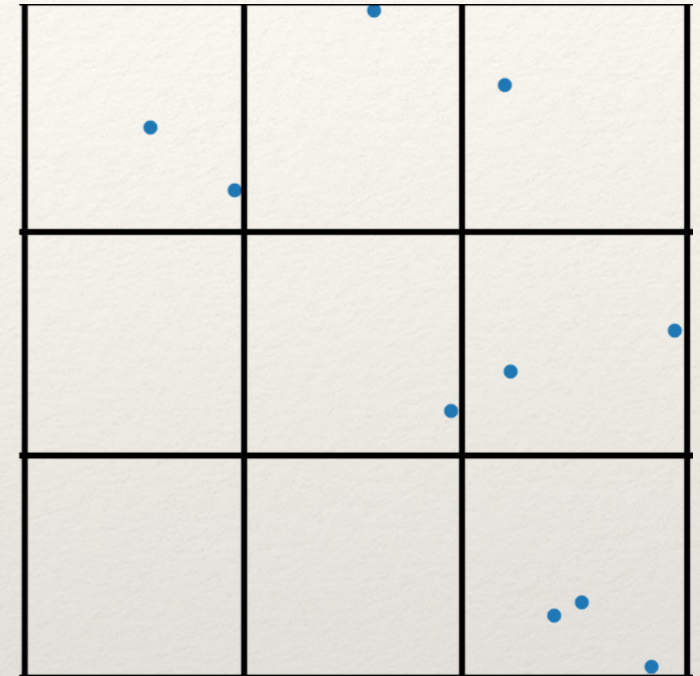
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# More general approach

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Inhomogeneous Poisson process for a given binning scheme:

$$p(\{\theta\}|\Lambda) = \prod_{k=1}^{N_{\text{bins}}} \frac{N_k(\Lambda)^{n_k} \exp[-N_k(\Lambda)]}{n_k!}$$



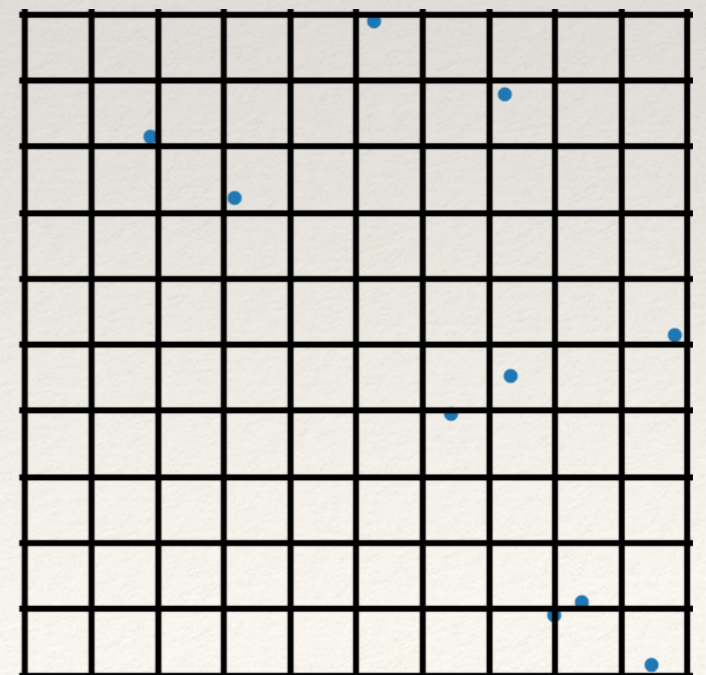
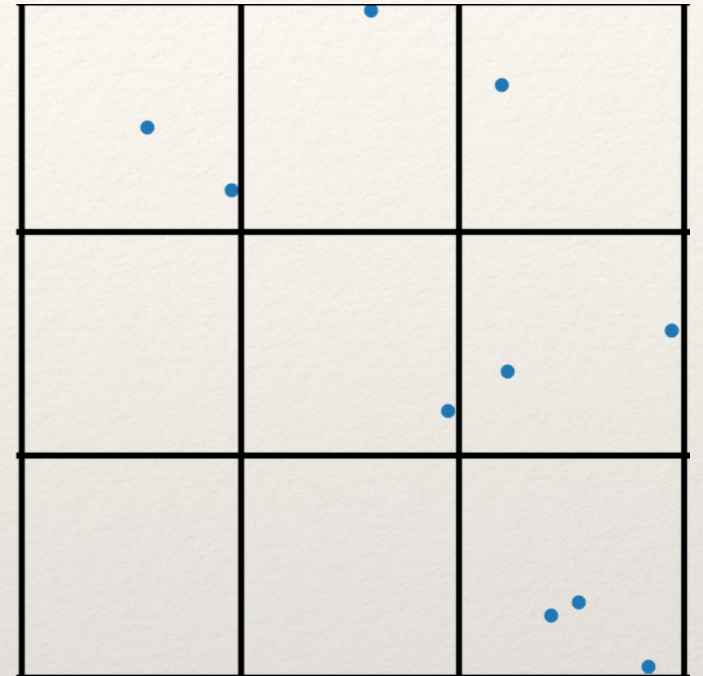
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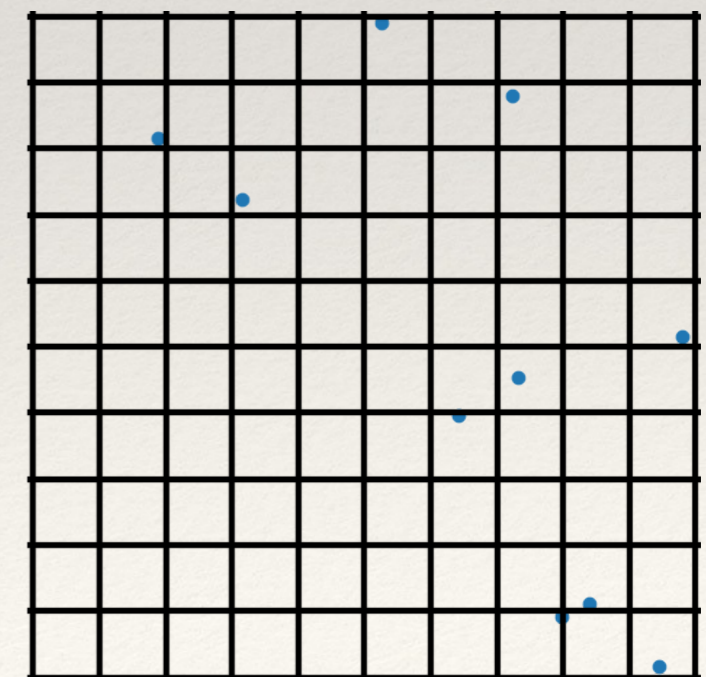
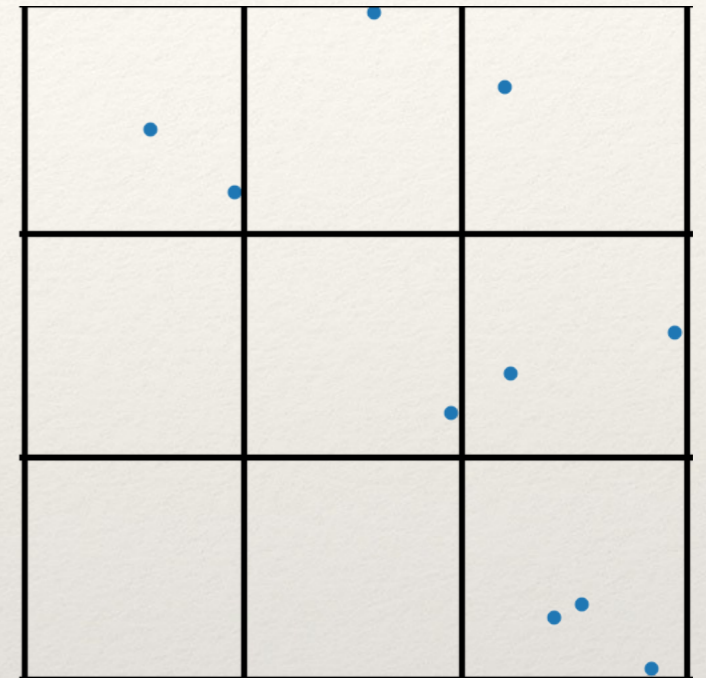
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$$p(\{\theta\} | \Lambda) = \exp[-N(\Lambda)] \prod_{k=1}^{N_{\text{obs}}} \frac{dN}{d\theta_k}(\Lambda)$$

$$N(\Lambda) = \int \frac{dN}{d\theta}(\Lambda) d\theta \quad \frac{dN}{d\theta} = N(\Lambda) p(\theta|\Lambda)$$



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# More general approach

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- ❖ Joint likelihood for all data/parameter pairs, including both **detected** events (indexed by  $i$ ) and **undetected** events (indexed by  $j$ )

$$p(\{x_i\}, \{\theta_i\}, \{x_j\}, \{\theta_j\} | \Lambda) = \left[ \prod_i^{N_{\text{obs}}} p(x_i | \theta_i) \frac{dN}{d\theta_i}(\Lambda) \right] \left[ \prod_j^{N_{\text{nobs}}} p(x_j | \theta_j) \frac{dN}{d\theta_j}(\Lambda) \right] \exp[-N(\Lambda)]$$

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- ❖ Marginalising over the unobserved data:

$$p(\{x_i\}, \{\theta_i\} | \Lambda) = \left[ \prod_{i=1}^{N_{\text{obs}}} p(x_i | \theta_i) \frac{dN}{d\theta_i}(\Lambda) \right] e^{-N(\Lambda)} \frac{N_{\text{ndet}}(\Lambda)^{N_{\text{nobs}}}}{N_{\text{nobs}}!}.$$

$$N_{\text{ndet}}(\Lambda) = \int_{F(x) < \text{threshold}} \int p(x | \theta) \frac{dN}{d\theta}(\Lambda) d\theta dd = N(\Lambda)(1 - \xi(\Lambda))$$

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- ❖ Marginalising over the parameters of observed events:

$$p(\{x\} | \Lambda) = e^{-N_{\text{det}}(\Lambda)} \prod_{i=1}^{N_{\text{obs}}} \int p(x_i | \theta) \frac{dN}{d\theta}(\Lambda) d\theta$$

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# More general approach

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- ❖ Marginalising over the unknown number of unobserved events:

$$p(\{x_i\}, \{\theta_i\} | \Lambda) = e^{-N(\Lambda)} e^{N_{\text{ndet}}(\Lambda)} \prod_{i=1}^{N_{\text{obs}}} p(x_i | \theta_i) \frac{dN}{d\theta_i}(\Lambda)$$

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**The selection function comes from marginalising over datasets with no detected events (and the parameters of undetected events)**

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# Case of backgrounds

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- ❖ Effectively, one dataset and no selection

$$p(x|\Lambda) = \int p(x|\Sigma)p(\Sigma|\Lambda)d\theta$$

With:

$$p(\Sigma|\Lambda) = \sum_{n=0}^{+\infty} \int d\{\theta\}_n p(\Sigma|\{\theta\}_n, n) \prod_{j=1}^n p(\theta_j|\Lambda) \frac{N(\Lambda)^n}{n!} e^{-n}$$

Ex:  $p(\Sigma|\{\theta\}_n, n) = \delta(\Sigma^2 - \sum_{i=1}^n |h(\theta_i)|^2)$



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Ex:  $p(\Sigma|\{\theta\}_n, n) = \delta(\Sigma^2 - \sum_{i=1}^n |h(\theta_i)|^2)$



- ❖ **Background + resolved:**

- Callister et al. APJL 2020: multiply likelihoods, okay for current LVK
- Toubiana, Gair 2601.04168: general case, including LISA Galactic binaries

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# Implementation

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- ❖ Use Bayes theorem for each event:  
$$p(\{x\}|\Lambda) = e^{-N_{\text{det}}(\Lambda)} \left( \prod_{i=1}^{N_{\text{obs}}} p(x_i) \right) \left( \prod_{i=1}^{N_{\text{obs}}} \int \frac{p(\theta|x_i)}{\pi_{\text{PE}}(\theta)} \frac{dN}{d\theta}(\Lambda) d\theta \right)$$

---

# Implementation

---

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❖ Monte-Carlo sampling

$$p(\{x\}|\Lambda) \propto e^{-N_{\text{det}}(\Lambda)} \prod_{i=1}^{N_{\text{obs}}} \frac{1}{N_s} \sum_{\theta \sim p(\theta|x_i)} \frac{1}{\pi_{\text{PE}}(\theta)} \frac{dN}{d\theta}(\Lambda) d\theta$$

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# Implementation

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Prior reweighting

# Implementation

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$$p(\{x\}|\Lambda) = e^{-N_{\text{det}}(\Lambda)} \left( \prod_{i=1}^{N_{\text{obs}}} p(x_i) \right) \left( \prod_{i=1}^{N_{\text{obs}}} \int \frac{p(\theta|x_i)}{\pi_{\text{PE}}(\theta)} \frac{dN}{d\theta}(\Lambda) d\theta \right)$$

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$$p(\{x\}|\Lambda) \propto e^{-N_{\text{det}}(\Lambda)} \prod_{i=1}^{N_{\text{obs}}} \frac{1}{N_s} \sum_{\theta \sim p(\theta|x_i)} \frac{1}{\pi_{\text{PE}}(\theta)} \frac{dN}{d\theta}(\Lambda) d\theta$$

Prior reweighting

❖ Similarly for selection function:

$$N_{\text{det}}(\Lambda) = \int_{F(x) > \text{threshold}} \int p(x|\theta) \frac{p(\theta|\Lambda)}{\pi_{\text{inj}}(\theta)} \pi_{\text{inj}}(\theta) d\theta dx$$

$$N_{\text{det}}(\Lambda) \simeq \frac{1}{N_{\text{inj}}} \sum_{\theta \sim \pi_{\text{inj}}(\theta), \text{det}} \frac{1}{\pi_{\text{inj}}(\theta)} \frac{dN}{d\theta}(\Lambda)$$

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# Implementation

---

❖ Use Bayes theorem for each event:

$$p(\{x\}|\Lambda) = e^{-N_{\text{det}}(\Lambda)} \left( \prod_{i=1}^{N_{\text{obs}}} p(x_i) \right) \left( \prod_{i=1}^{N_{\text{obs}}} \int \frac{p(\theta|x_i)}{\pi_{\text{PE}}(\theta)} \frac{dN}{d\theta}(\Lambda) d\theta \right)$$

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$$N_{\text{det}}(\Lambda) = \int_{F(x) > \text{threshold}} \int p(x|\theta) \frac{p(\theta|\Lambda)}{\pi_{\text{inj}}(\theta)} \pi_{\text{inj}}(\theta) d\theta dx$$

$$N_{\text{det}}(\Lambda) \simeq \frac{1}{N_{\text{inj}}} \sum_{\theta \sim \pi_{\text{inj}}(\theta), \text{det}} \frac{1}{\pi_{\text{inj}}(\theta)} \frac{dN}{d\theta}(\Lambda)$$

**Take care with efficiency of Monte-Carlo estimators!**

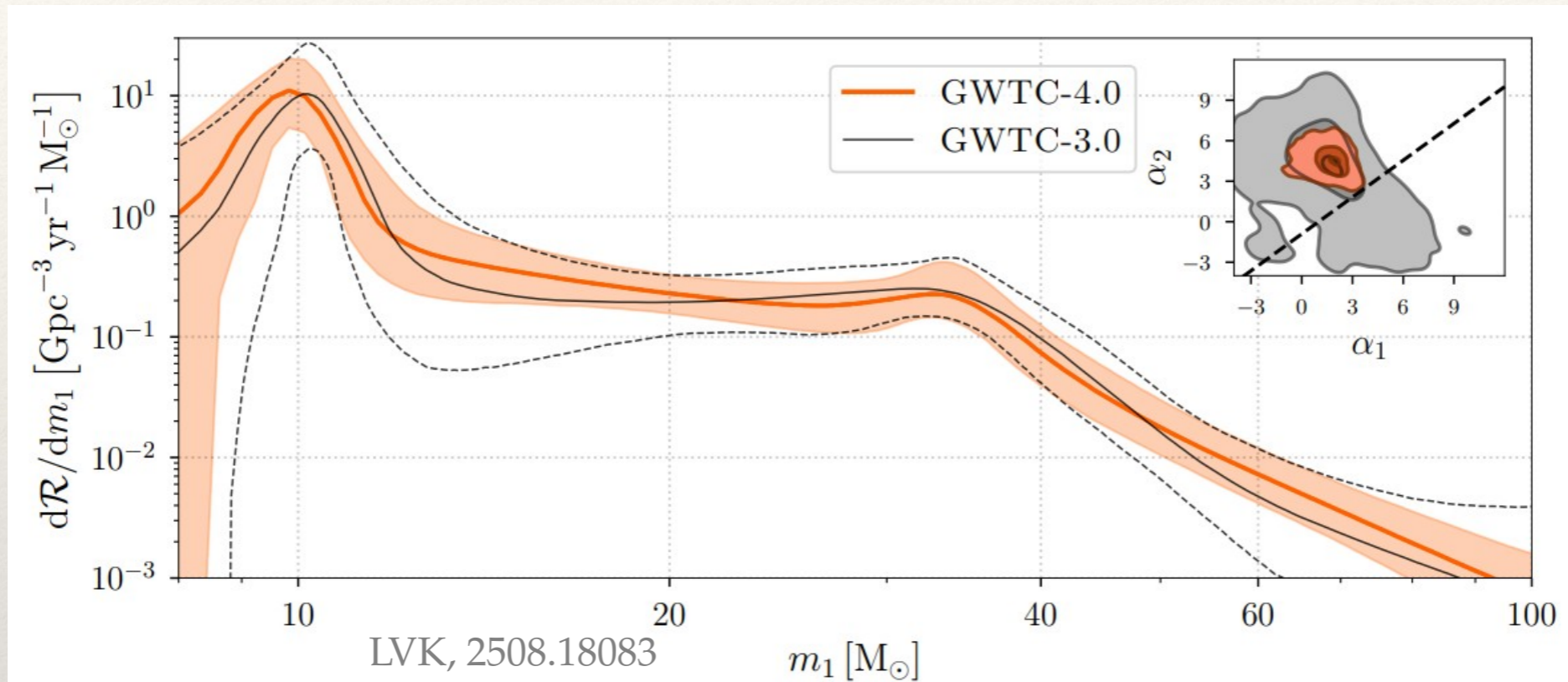
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# Astrophysical inference on GW data

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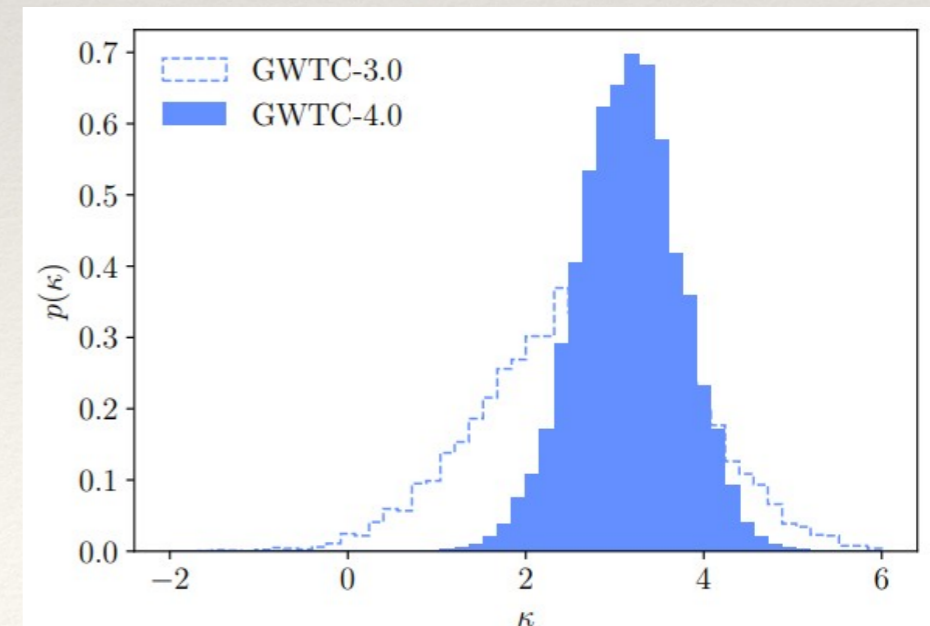
- ❖ Assume a form for  $p(\theta|\Lambda)$   
Can be:
  - Astrophysical
  - parametric
  - non-parametric

# Parametric model (LVK)



Broken power-law + 2 peaks

- ❖ Peaks at  $\sim 10M_{\odot}$  and  $\sim 35M_{\odot}$ , steeper power-law at high mass
- ❖ Evidence for rate evolution, with the rate higher in the past.  $R(z) = R_0(1+z)^{\kappa}$



# Parametric model 2D (LVK)

**Linear:**

$$p(\chi_{\text{eff}}|q) \propto \exp\left[-\frac{(\chi_{\text{eff}} - \mu(q))^2}{2\sigma(q)^2}\right]$$

$$\mu(q) = \mu_0 + \alpha(q - 1)$$

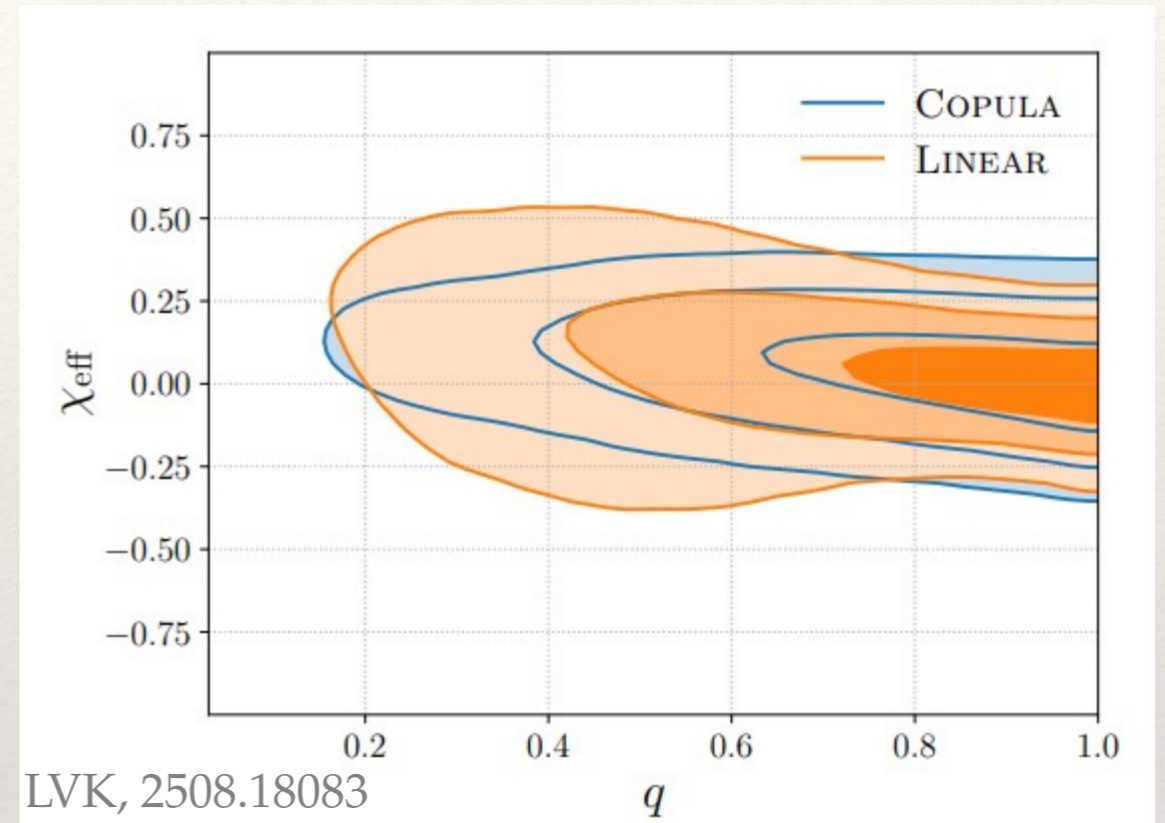
$$\log_{10} \sigma(q) = \log_{10}(\sigma_0) + \beta(q - 1)$$

**Copula:**

$$p(\chi_{\text{eff}}, q) = p(\chi_{\text{eff}})p(q)C(\chi_{\text{eff}}, q)$$

$C(\chi_{\text{eff}}, q)$  preserves the marginal

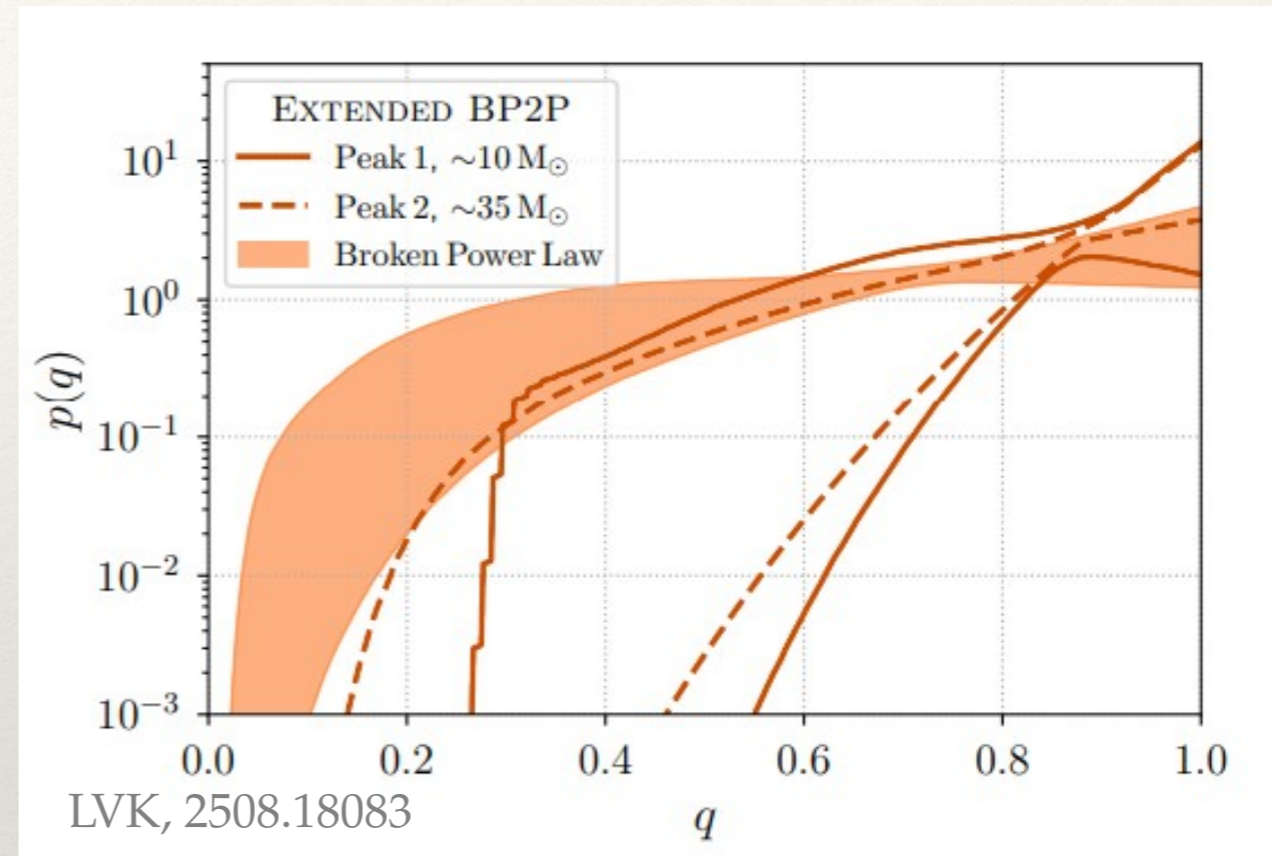
- ❖ Weaker evidence in GWTC-4 compared to GWTC-3



# Parametric model nD (LVK)

$$p(m_1, q|\Lambda) = \sum_{i=1}^n \beta_i p_i(m_1|\Lambda) p_i(q|\Lambda)$$

- ❖ Black holes in the  $10M_{\odot}$  peak may favour unequal mass ratios

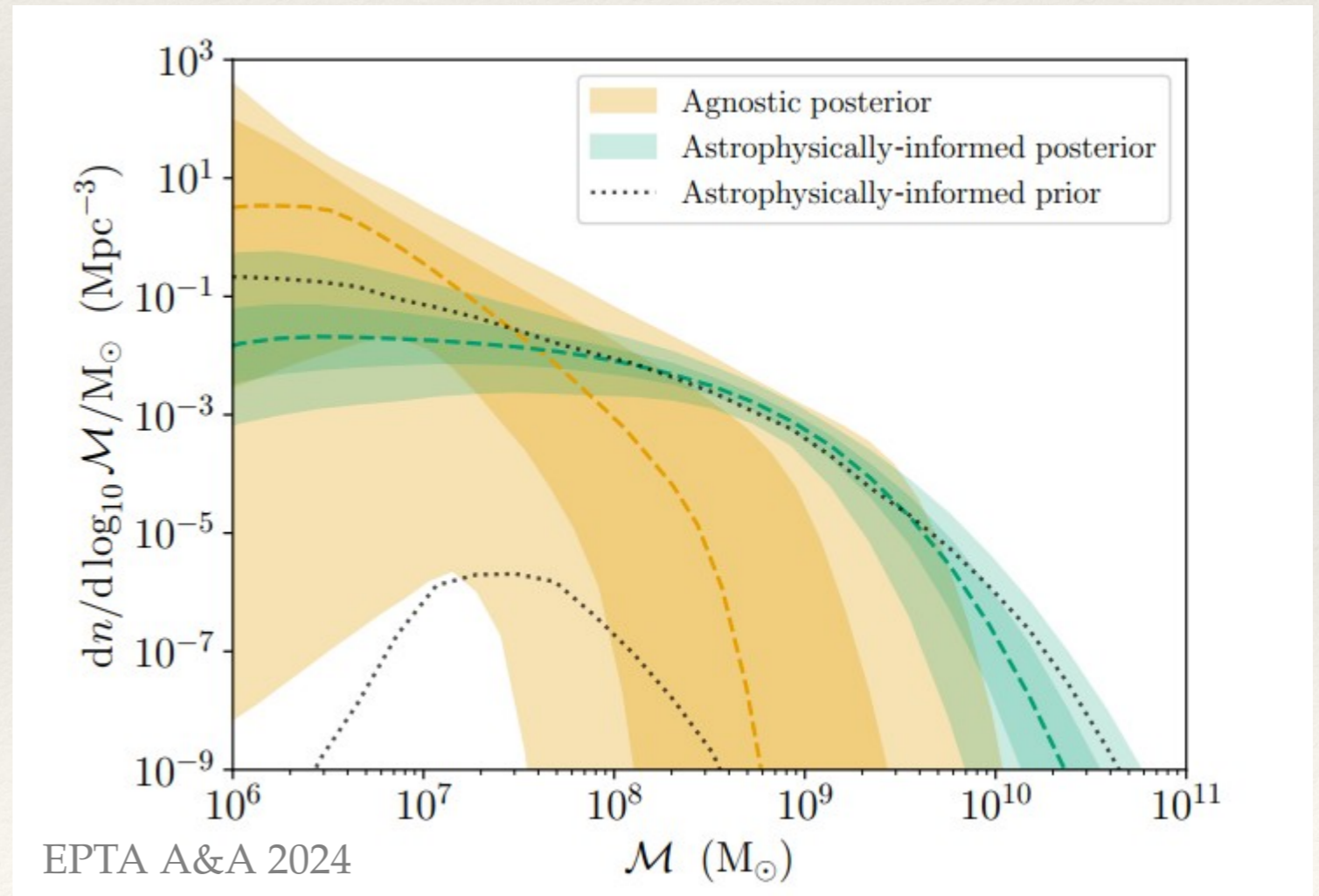


- ❖ Number of components can be inferred with reversible-jump MCMC (Cheng et al. to appear)

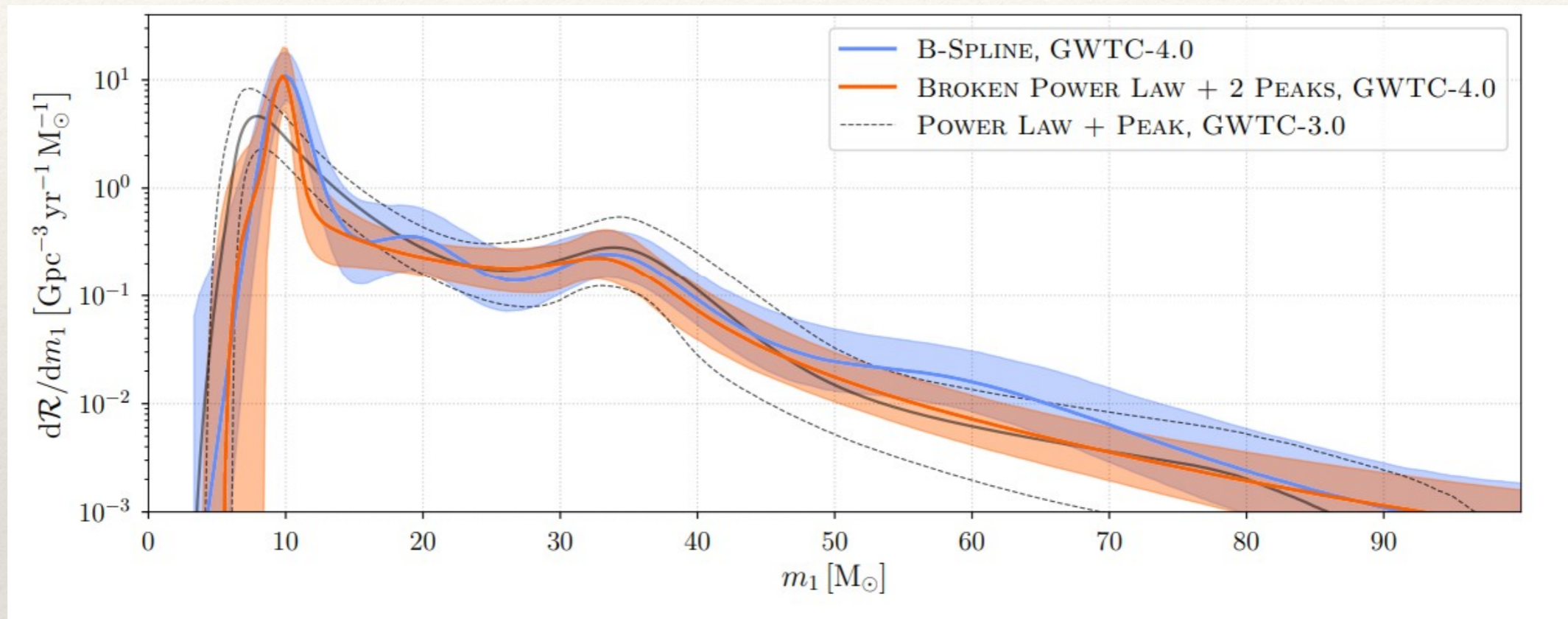
# Parametric model (PTA)

$$\frac{d^2 n}{dz d \log_{10} \mathcal{M}} = \dot{n}_0 \left[ \left( \frac{\mathcal{M}}{10^7 M_\odot} \right)^{-\alpha_{\mathcal{M}}} e^{-\mathcal{M}/\mathcal{M}_*} \right] \left[ (1+z)^{\beta_z} e^{-z/z_0} \right] \frac{dt_R}{dz}$$

$$\mathcal{M} = \left( \frac{m_1^3 m_2^3}{m_1 + m_2} \right)^{1/5}$$



# Non-parametric model (LVK)

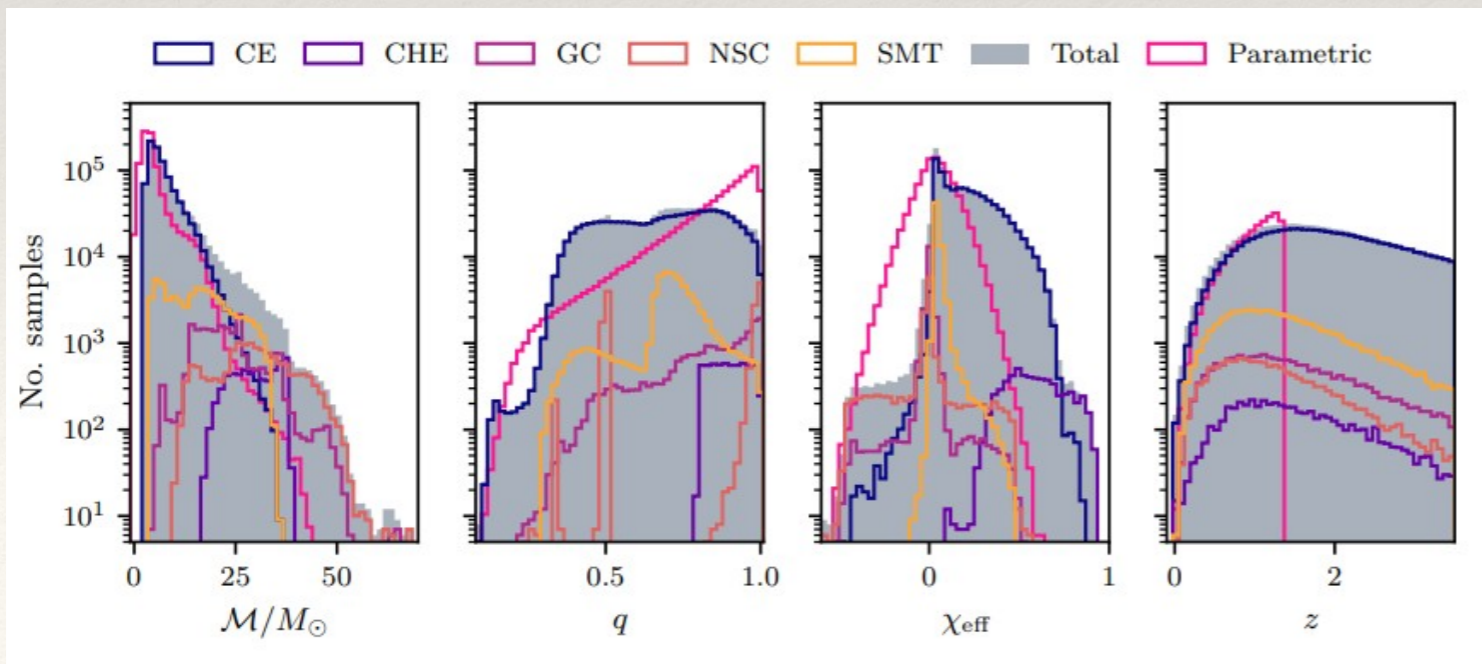
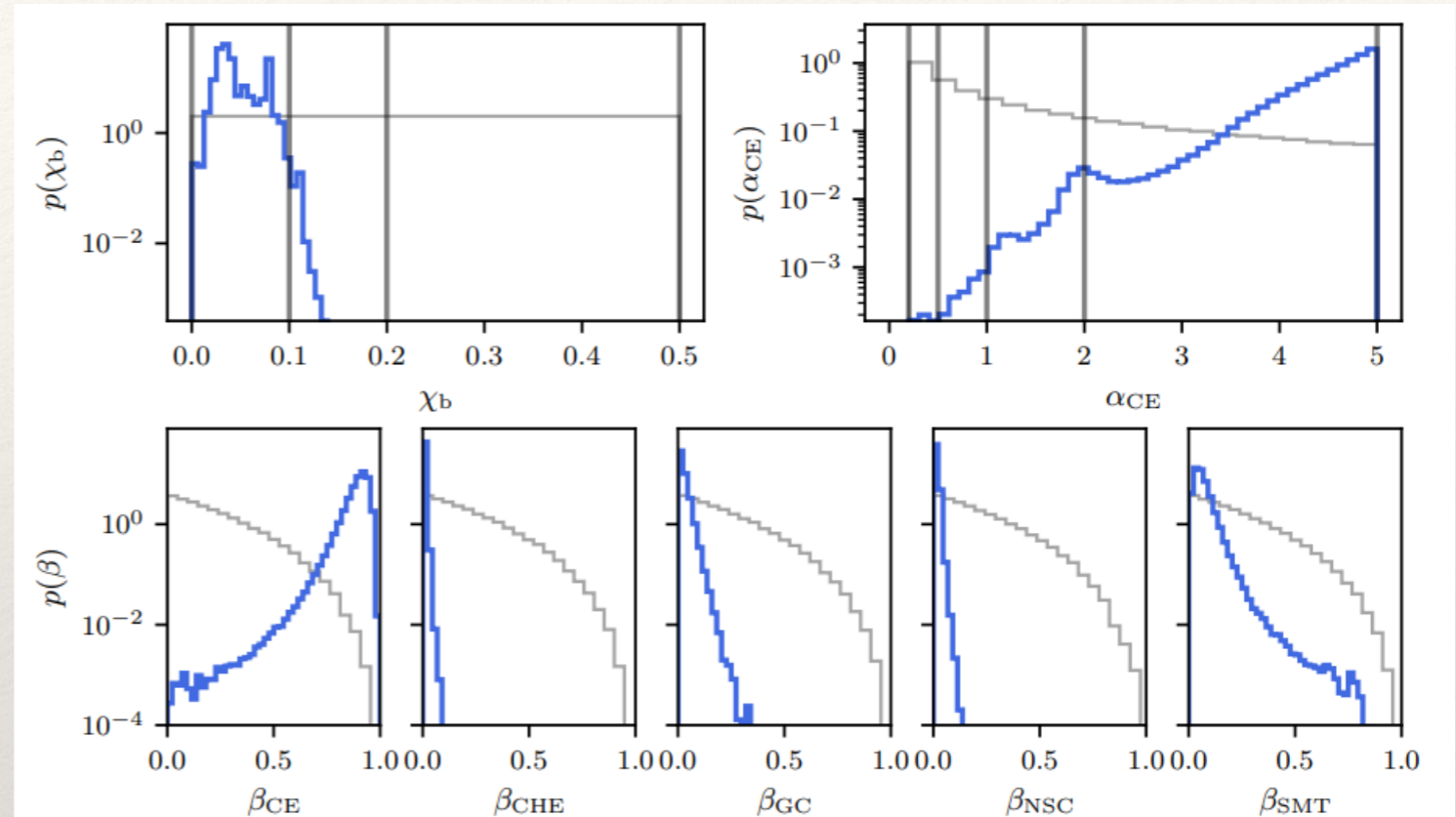


- ❖ Additional peak at  $\sim 20M_{\odot}$ , excess at high mass?
- ❖ See also Toubiana et al. MNRAS 2023 for optimisation of number of knots with reversible-jump MCMC

# Astrophysical (LVK)

$$p(\theta|\Lambda) = \sum_{i=1}^5 \beta_i p_i(\theta|\Lambda)$$

- ❖ CE: common envelope
- ❖ CHE: chemically homogeneous evolution
- ❖ GC: globular cluster
- ❖ NSC: nuclear stellar cluster
- ❖ SMT: stable mass transfer



Colloms et al. A&A 2025

Favours formation through common envelope

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# Astrophysical inference on GW data

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## Astrophysical

### ❖ Pros:

Gives direct information on astrophysical processes

### ❖ Cons:

Huge uncertainty on astro models. We might not include all channels (see Cheng et al., 2307.03129, Raikman et al., 2310.10736)

Requires some way to evaluate pdf from samples and to interpolate (see Toubiana et al., PRD 2021 for systematic errors)

## Parametric

### ❖ Pros:

Analytic pdfs, easy to evaluate  
Some astrophysical meaning

### ❖ Cons:

Little flexibility  
Not so much astrophysical meaning

## Non-parametric

### ❖ Pros:

Very flexible  
Requires less a priori knowledge

### ❖ Cons:

Parameters have no astrophysical meaning  
Complexity is a priori arbitrary (might be alleviated using reversible-jump MCMC)

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# “Mock LVK” toy model

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- ❖ Mass distributed according to power-law+peak
- ❖ Known selection function
- ❖ Gaussian error on each measurement
- ❖ 100 detected events

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# Posterior predictive checks

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- ❖ Test if the assumed model is a good fit to the data

---

# Posterior predictive checks

---

❖ Test if the assumed model is a good fit to the data

❖ **Prior predictive distribution:**

$$p(x) = \int p(x | \Lambda) p(\Lambda) d\Lambda$$

Distribution of data within the model assumed in the prior

---

# Posterior predictive checks

---

- ❖ Test if the assumed model is a good fit to the data

- ❖ **Prior predictive distribution:**

$$p(x) = \int p(x | \Lambda) p(\Lambda) d\Lambda$$

Distribution of data within the model assumed in the prior

- ❖ **Posterior predictive distribution:**

$$p(x_{\text{new}} | \{x\}_{\text{old}}) = \int p(x_{\text{new}} | \Lambda) p(\Lambda | \{x\}_{\text{old}}) d\Lambda$$

---

# Posterior predictive checks

---

- ❖ Test if the assumed model is a good fit to the data

- ❖ **Prior predictive distribution:**

$$p(x) = \int p(x | \Lambda) p(\Lambda) d\Lambda$$

Distribution of data within the model assumed in the prior

- ❖ **Posterior predictive distribution:**

$$p(x_{\text{new}} | \{x\}_{\text{old}}) = \int p(x_{\text{new}} | \Lambda) p(\Lambda | \{x\}_{\text{old}}) d\Lambda$$

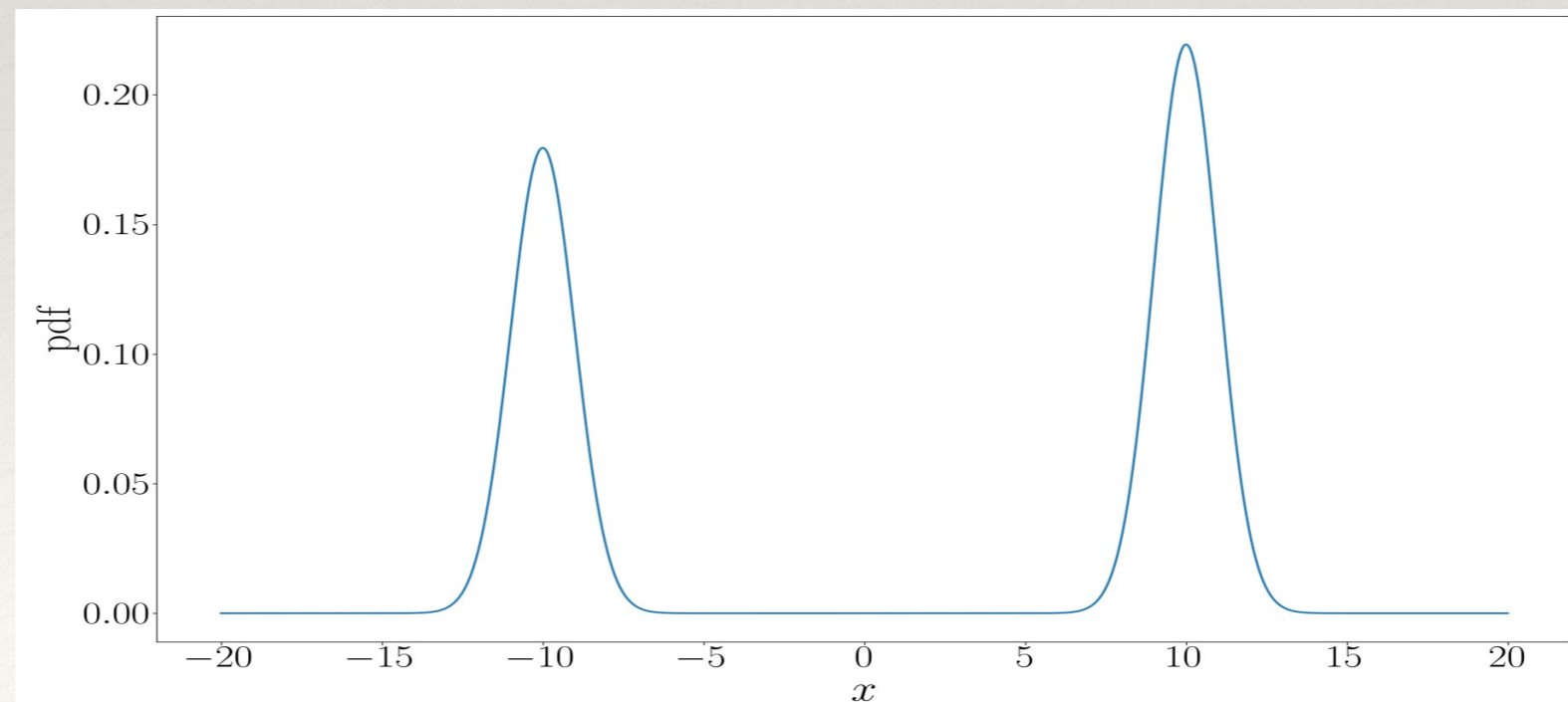
Distribution of new datasets based on the model fitted to the data.  
Observed data should lie within this distribution if model is good.

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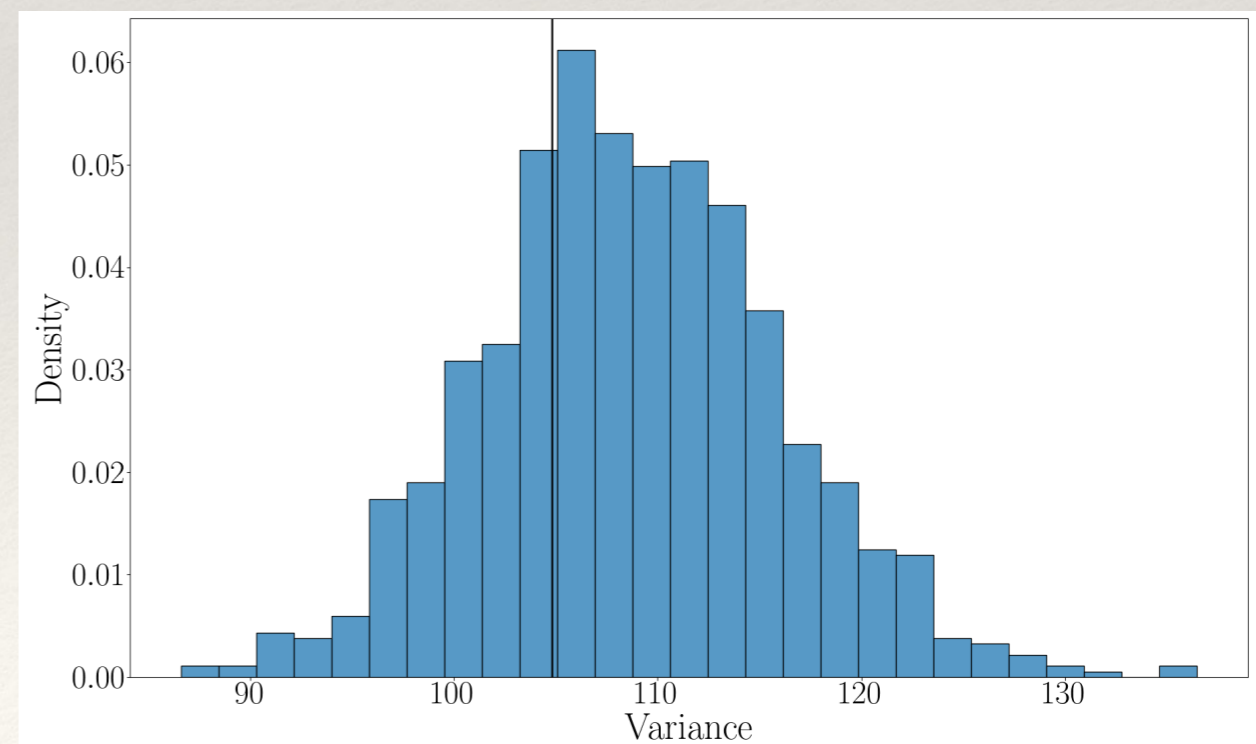
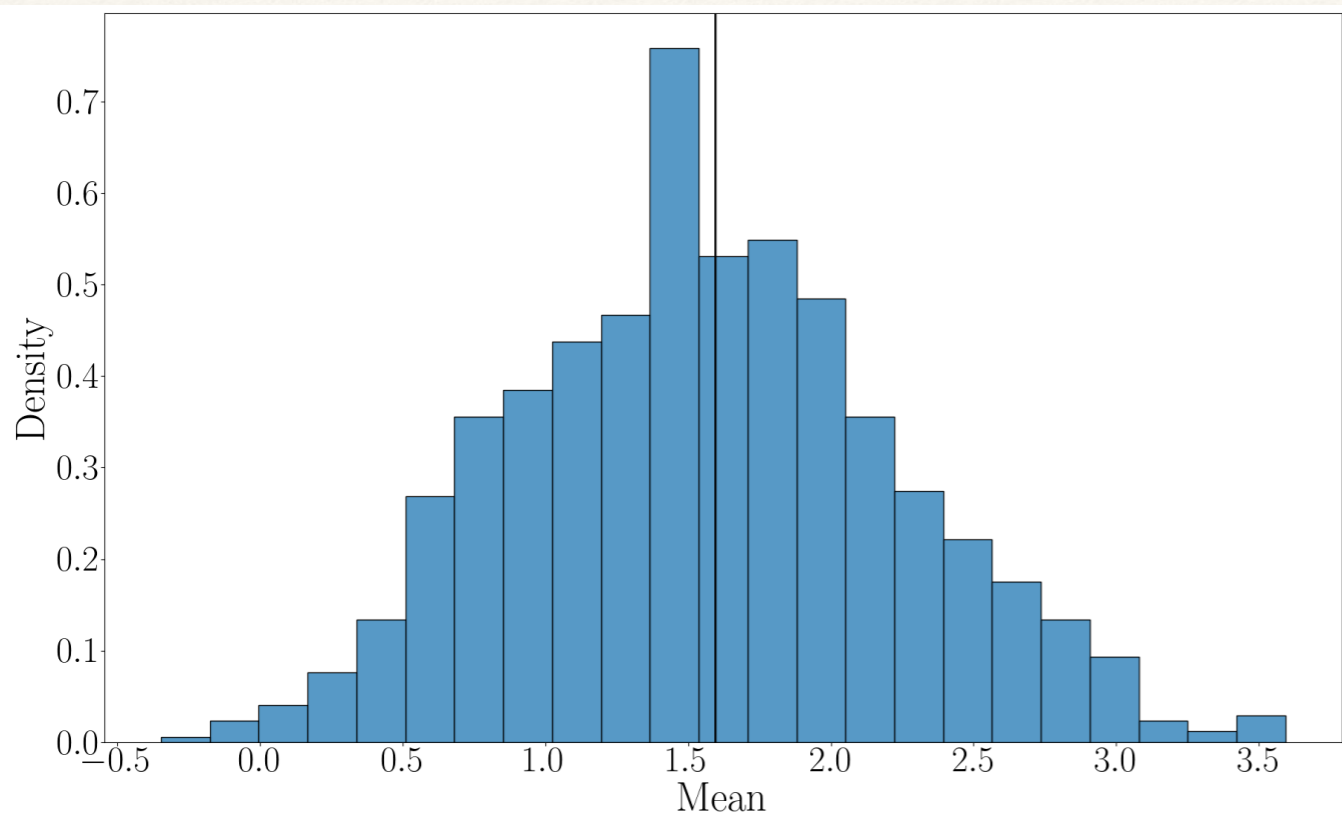
# Example: gaussian fit

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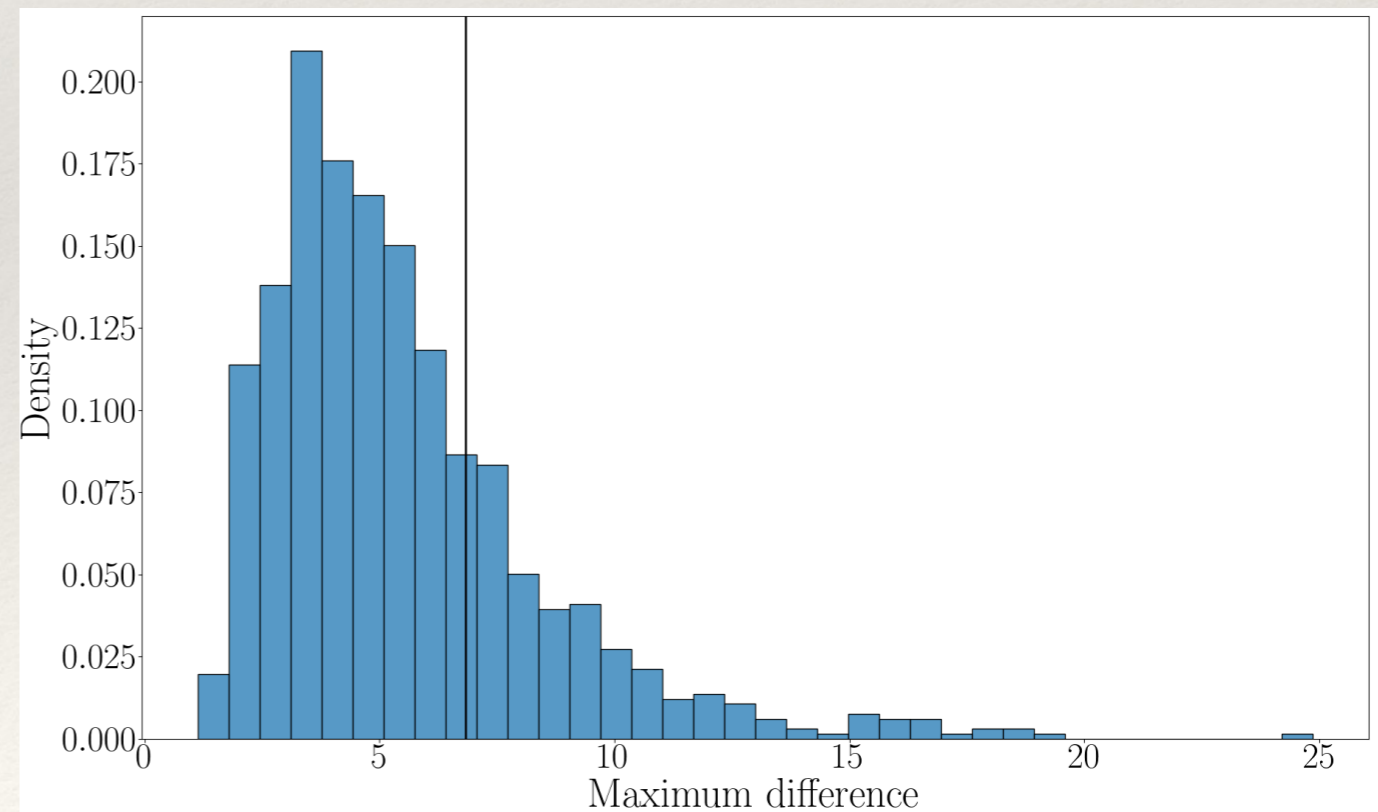
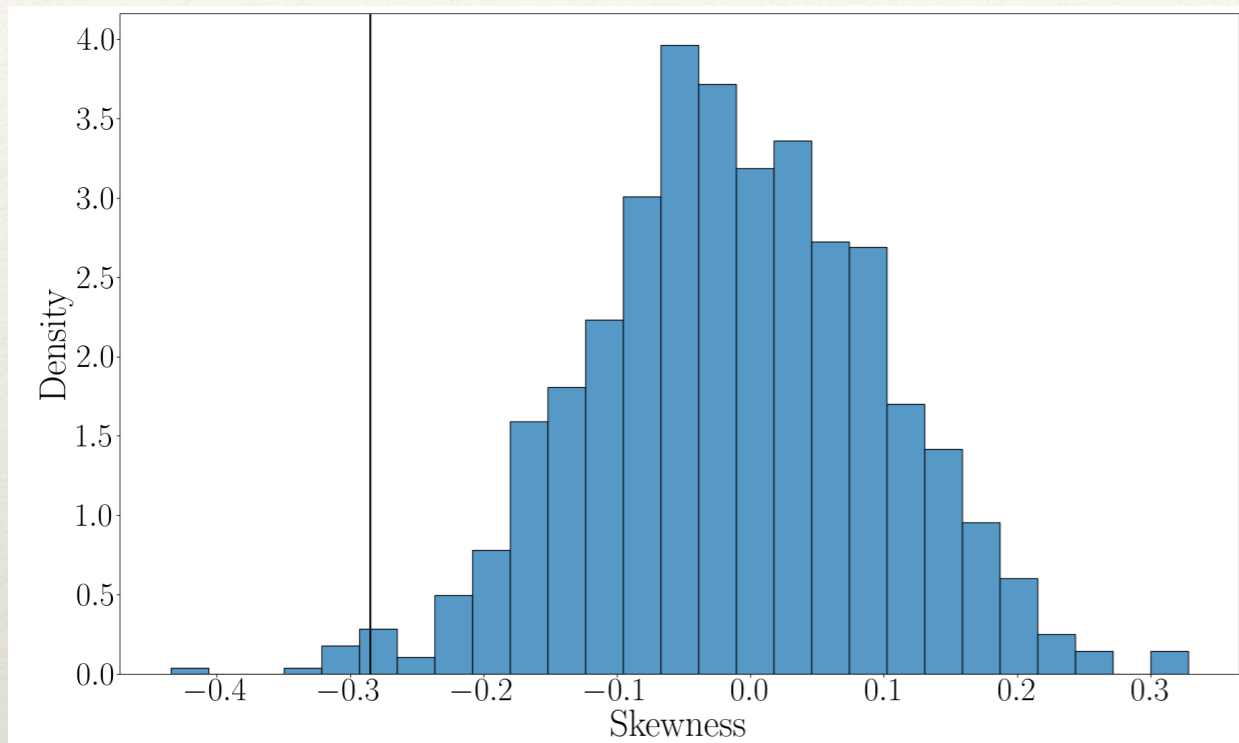
- ❖ The predictive distribution can be used to compute the distribution of summary quantities. The value of those summary quantities in the observed data can then be compared to these distributions.
- ❖ It is better to choose quantities that are somewhat “orthogonal” to what is adjusted to fit the data.
- ❖ Example: we try to fit a Gaussian to the following distribution:
- ❖ We assume a Gaussian measurement error of 2



# Example: gaussian fit



# Example: gaussian fit



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# Posterior predictive checks

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- ❖ Posterior predictive checks are “good practice”.
- ❖ Can help build intuition how to improve models.
- ❖ But are often computationally expensive...